

Cosmic Flows

Lecture course
University Groningen
Apr. 2009-July 2009

Gravitational Instability

$$\frac{\partial \delta}{\partial t} + \frac{1}{a} \nabla \cdot (1 + \delta) \mathbf{v} = 0$$

$$\frac{\partial \mathbf{v}}{\partial t} + \frac{\dot{a}}{a} \mathbf{v} + \frac{1}{a} (\mathbf{v} \cdot \nabla) \mathbf{v} = -\frac{1}{a} \nabla \phi$$

$$\nabla^2 \phi = 4\pi G \bar{\rho} a^2 \delta(\mathbf{x}, t)$$

Gravitational Instability

$$\frac{\partial \delta}{\partial t} + \frac{1}{a} \nabla \cdot \mathbf{v} = 0$$

$$\frac{\partial \mathbf{v}}{\partial t} + \frac{\dot{a}}{a} \mathbf{v} = -\frac{1}{a} \nabla \phi$$

$$\nabla^2 \phi = \frac{3}{2} \Omega H^2 a^2 \delta(\mathbf{x}, t)$$

Gravitational Instability

The linear system of structure growth equations can be written in terms of a second order differential equation,

$$\frac{\partial^2 \delta}{\partial t^2} + 2 \frac{\dot{a}}{a} \frac{\partial \delta}{\partial t} = \frac{3}{2} \Omega_0 H_0^2 \frac{1}{a^3} \delta$$

Gravitational Instability

... whose two solutions are separable in time and space,
leading to a universal "density growth factor" $D(t)$,

$$\delta(\mathbf{x}, t) = D_1(t) \Delta_1(\mathbf{x}) + D_2(t) \Delta_2(\mathbf{x})$$

"Growing Mode"

"Decaying Mode"

Linear Density Growth

... whose two solutions are separable in time and space,
leading to a universal "density growth factor" $D(t)$,

$$\delta(\mathbf{x}, t) = D_1(t) \Delta_1(\mathbf{x}) + D_2(t) \Delta_2(\mathbf{x})$$

"Growing Mode"

"Decaying Mode"

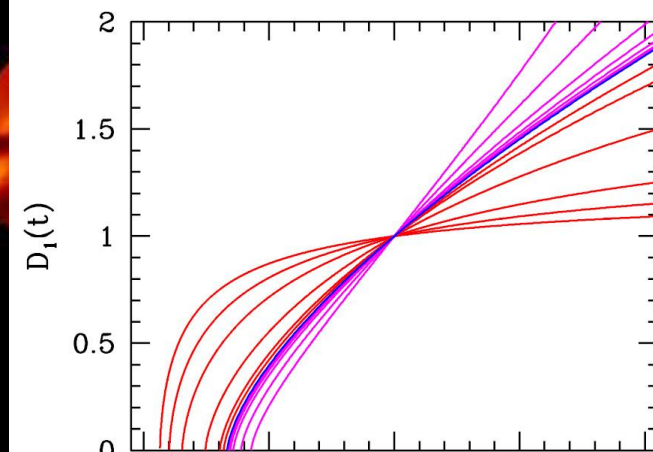
Linear Density Growth

... The universal "density growth factor" $D(t)$ can be computed for any cosmology through the integral

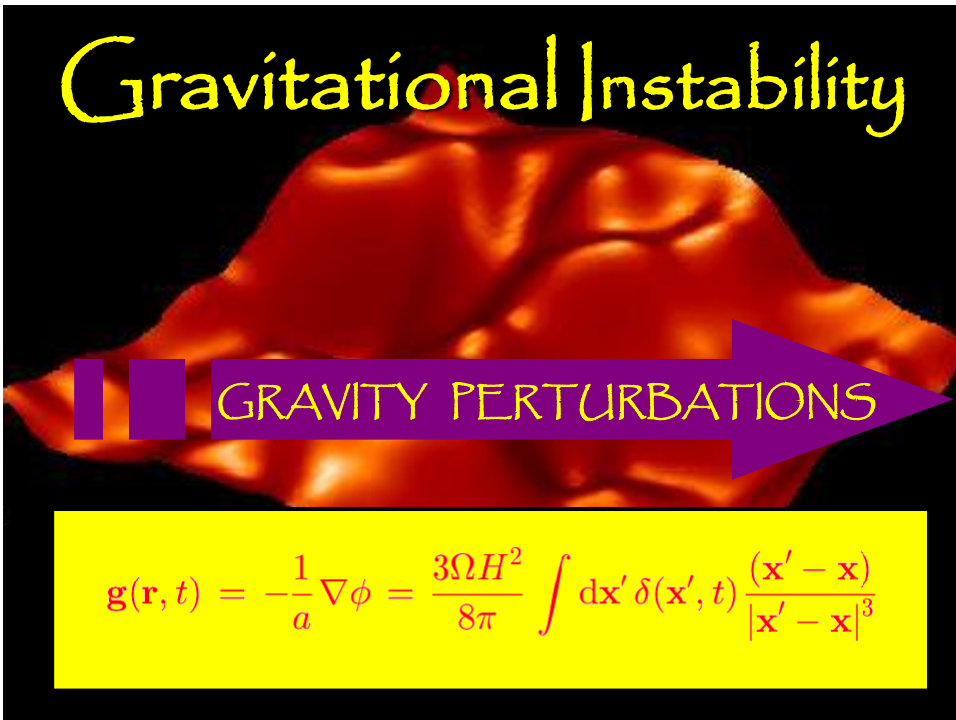
$$D(t) \approx H(t) \int \frac{dt}{a^2 H^2(t)}$$

Linear Density Growth

Linear Perturbation Evolution:
Density evolution: Growing Mode D_1

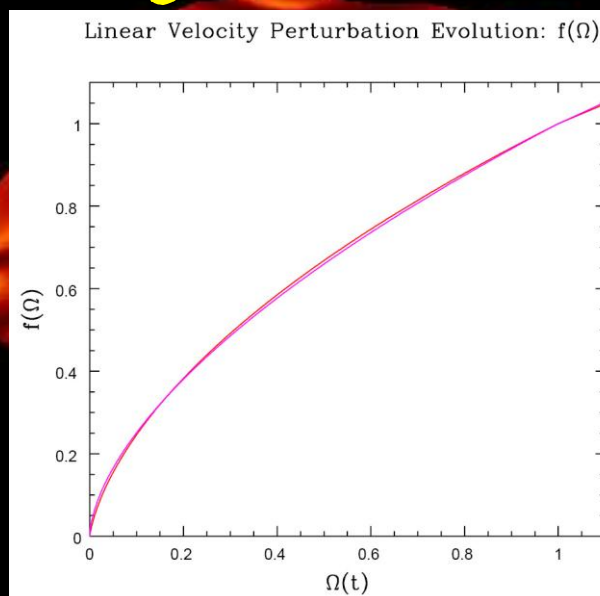


Gravitational Instability



$$\mathbf{g}(\mathbf{r}, t) = -\frac{1}{a} \nabla \phi = \frac{3\Omega H^2}{8\pi} \int d\mathbf{x}' \delta(\mathbf{x}', t) \frac{(\mathbf{x}' - \mathbf{x})}{|\mathbf{x}' - \mathbf{x}|^3}$$

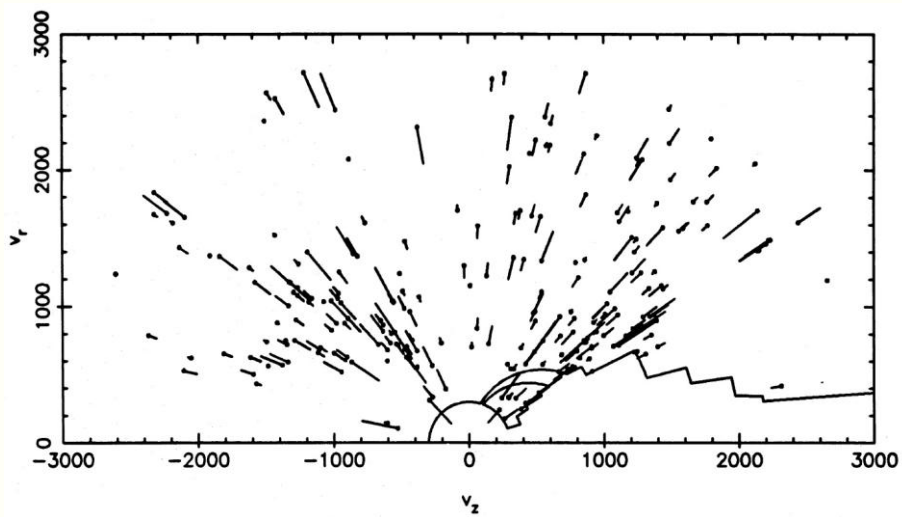
Velocity Perturbations



Local Universe Peculiar Velocities

Local Superclusterflow

Lilje, Yahil & Jones 1986

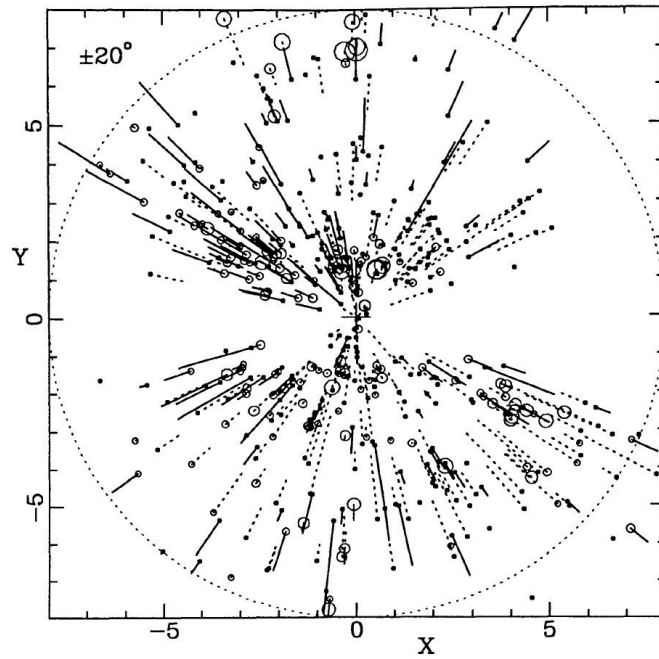


Mark III

Peculiar velocities
Local Universe

Willick, Strauss et al.

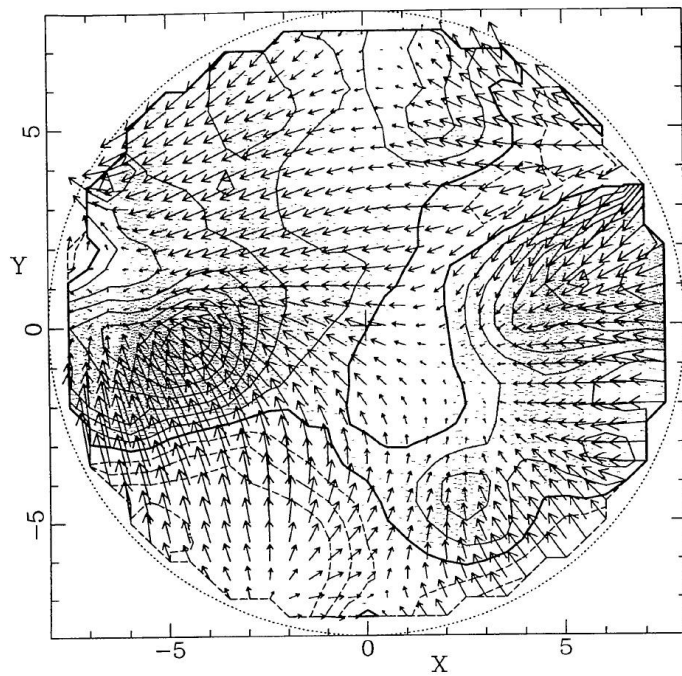
First flow maps:
7 Samurai

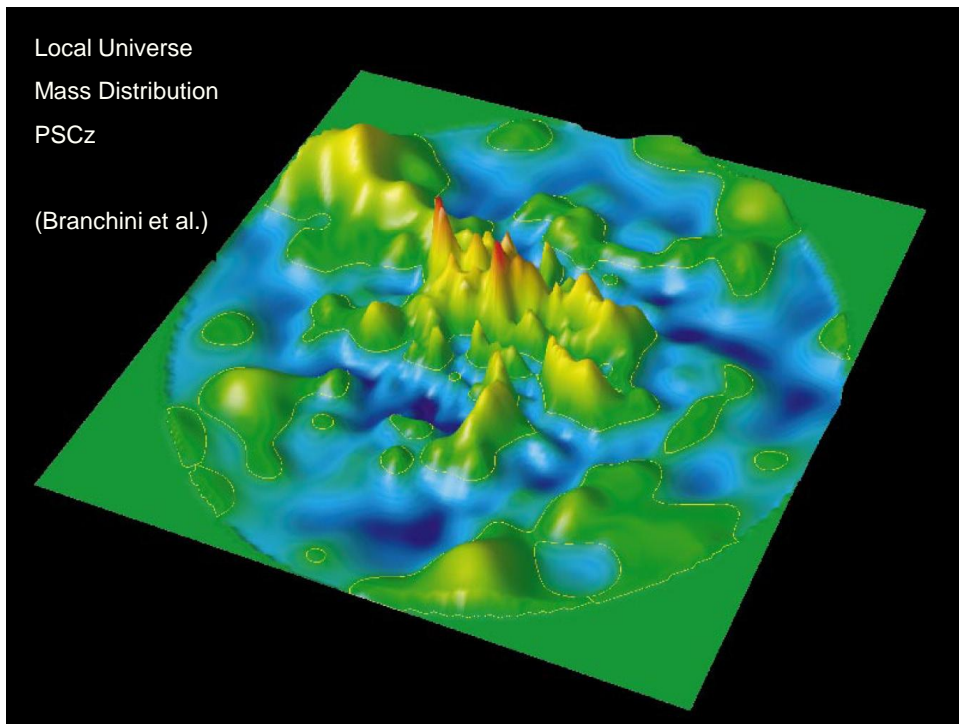
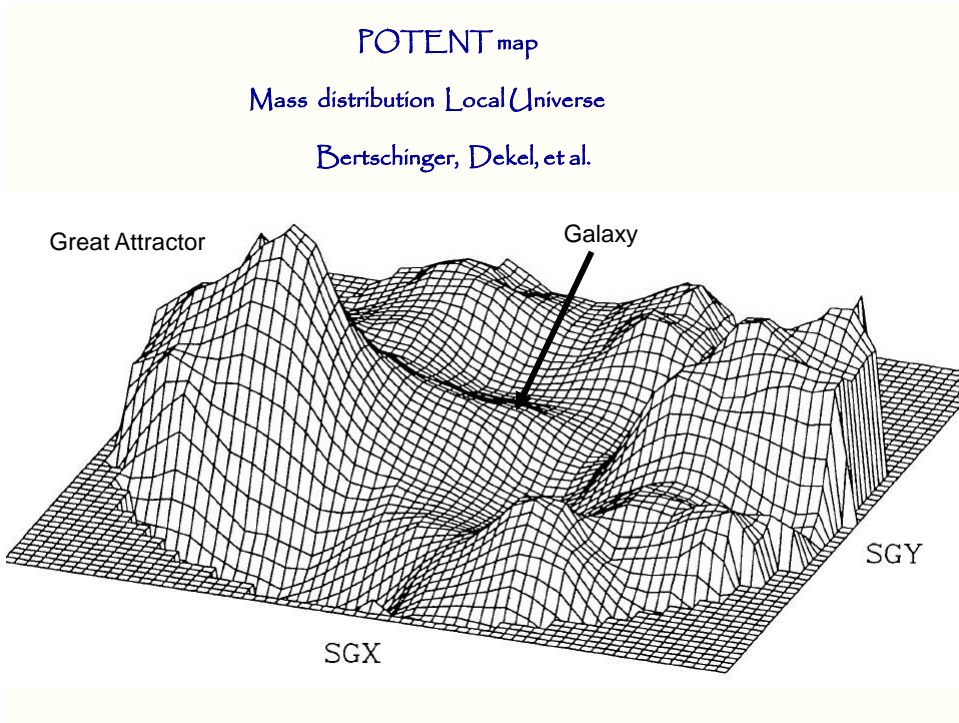


POTENT map

Mass distribution
Local Universe

Bertschinger,
Dekel, et al.





Great Attractor

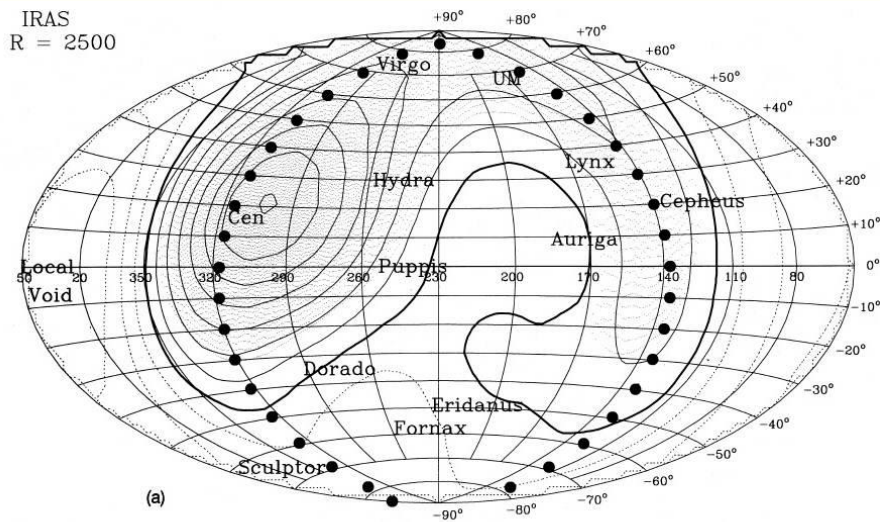
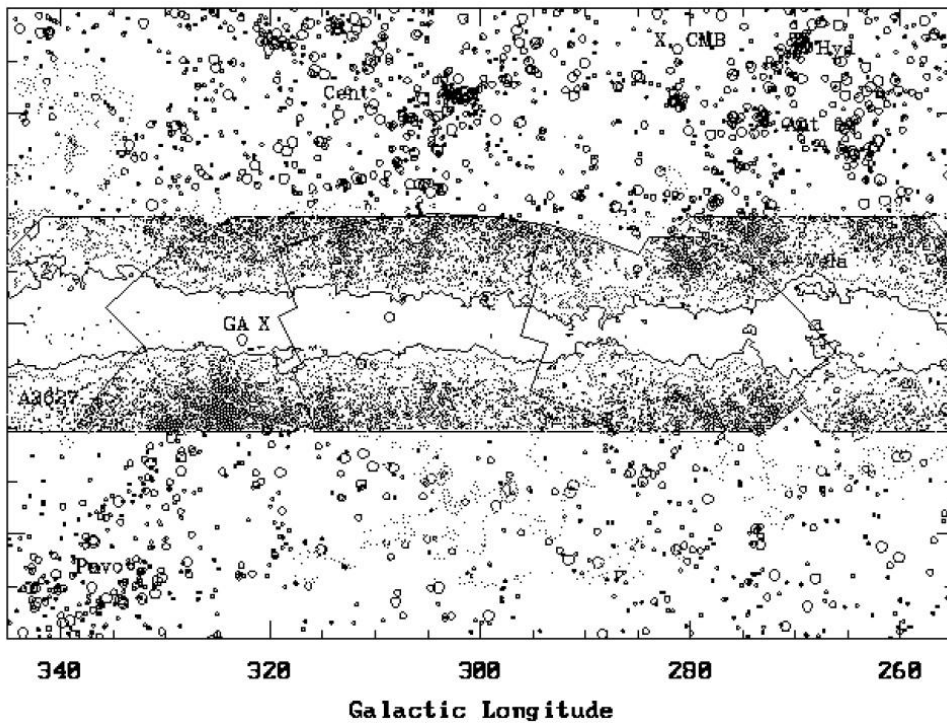
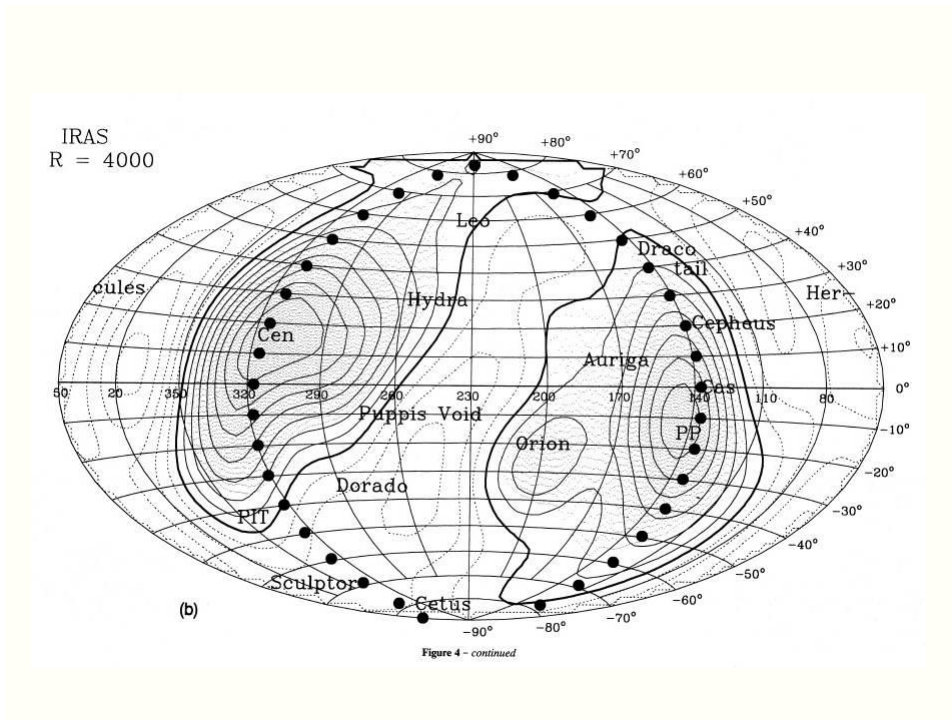
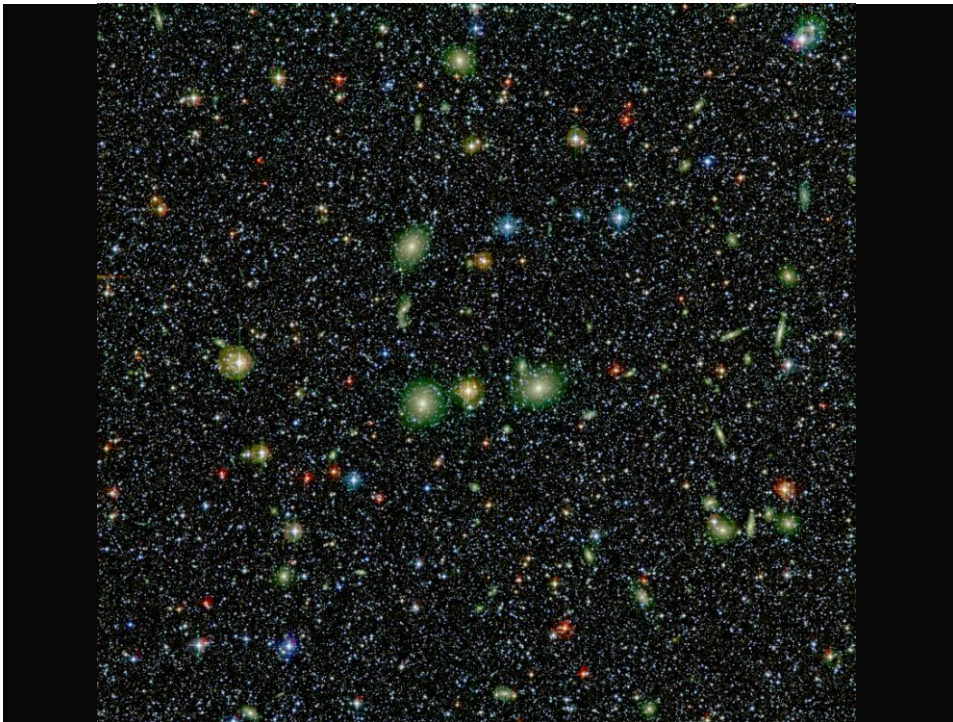
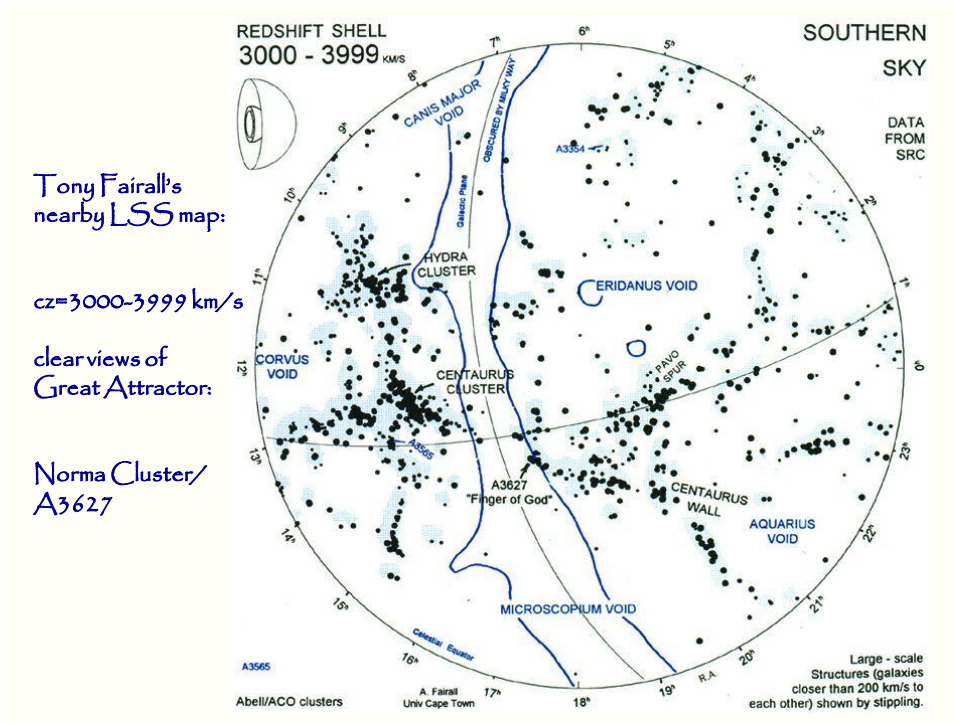


Figure 4. The galaxy density fluctuation field from the *IRAS* 1.9-Jy survey (by Yahil et al. 1991). Coordinates, smoothing, contours and shell distances are as in Fig. 3.





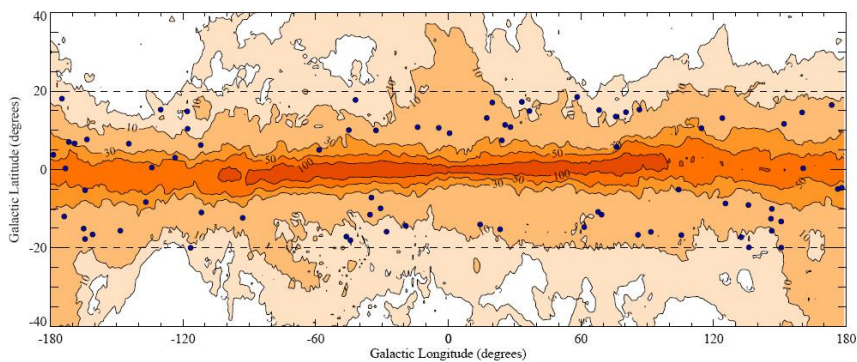


Fig. 16. Distribution in Galactic coordinates of the 76 by Ebeling et al. [129] so far spectroscopically confirmed X-ray clusters (solid dots) of which 80% were previously unknown. Superimposed are Galactic HI column densities in units of 10^{20} cm^{-2} (Dickey & Lockman 1990). Note that the region of relatively high absorption ($N_{\text{HI}} > 5 \times 10^{21} \text{ cm}^{-2}$) actually is very narrow and that clusters could be identified to very low latitudes

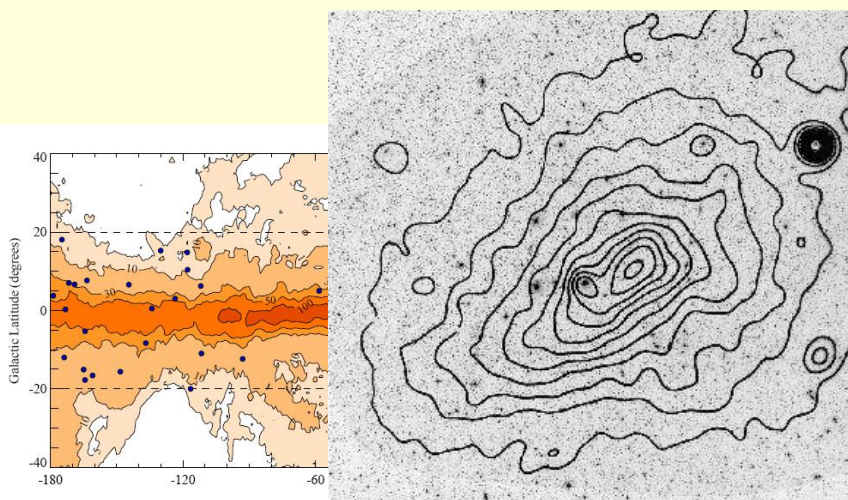
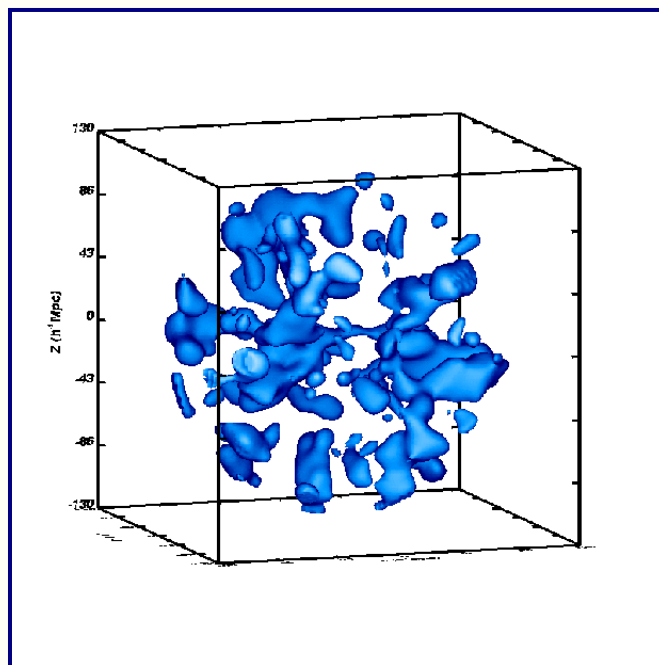
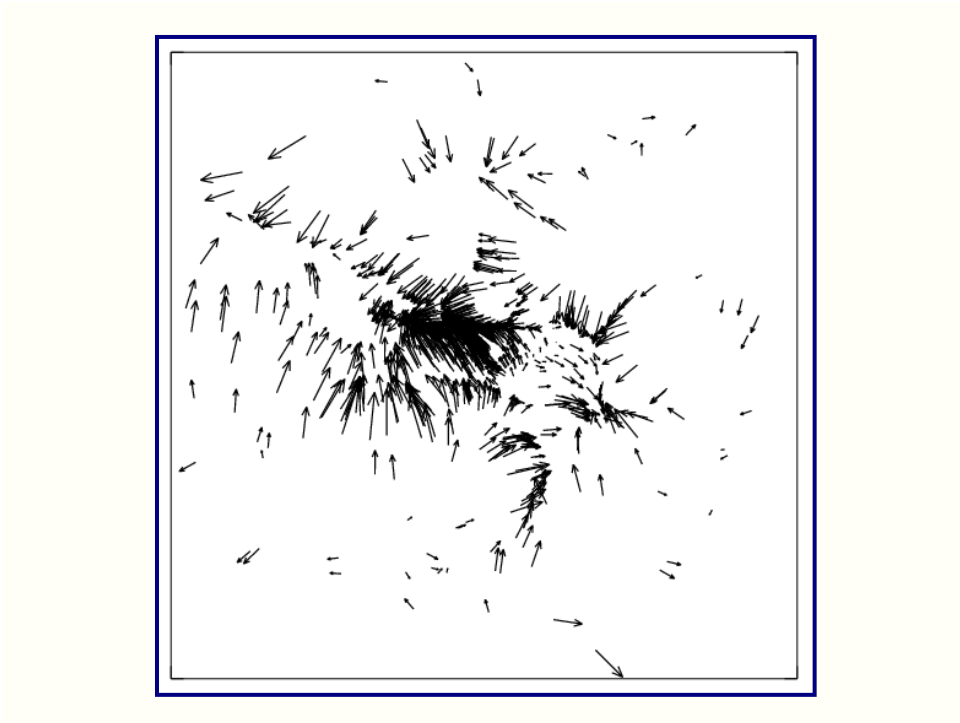
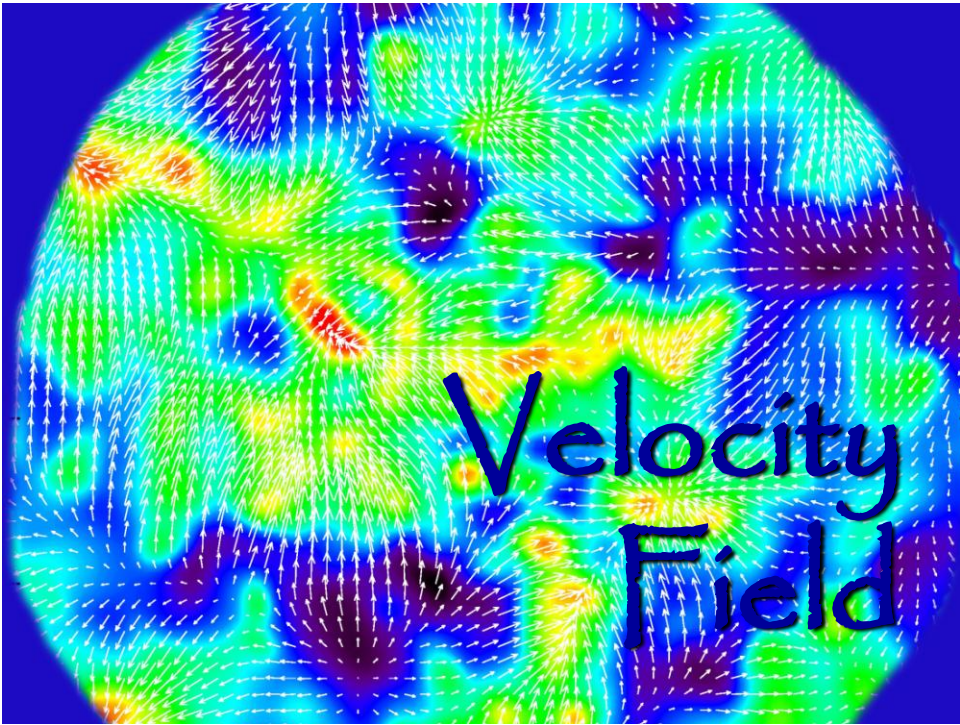


Fig. 16. Distribution in Galactic coordinates of the 76 by Ebeling et al. [129] so far spectroscopically confirmed X-ray clusters (solid dots) of which 80% were previously unknown. Superimposed are Galactic HI column densities in units of 10^{20} cm^{-2} (Dickey & Lockman 1990). Note that the region of relatively high absorption ($N_{\text{HI}} > 5 \times 10^{21} \text{ cm}^{-2}$) actually is very narrow and that clusters could be identified to very low latitudes

PSCz Survey

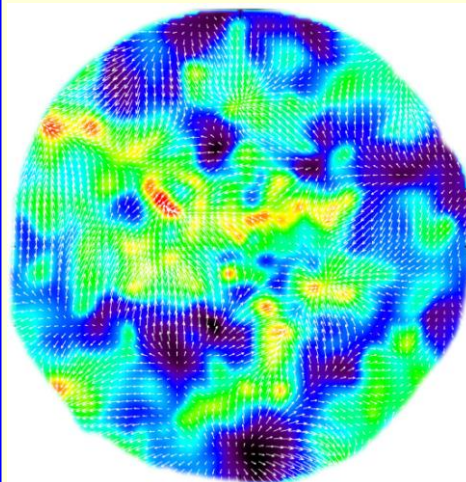
Dynamics:



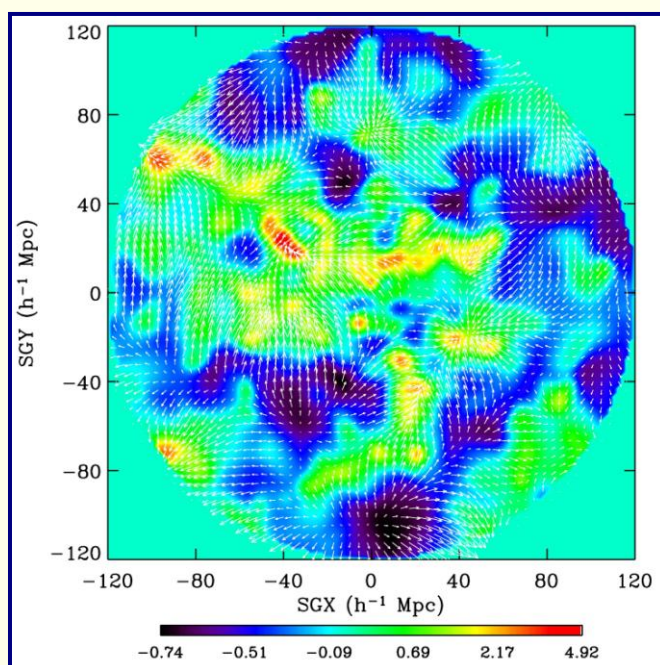


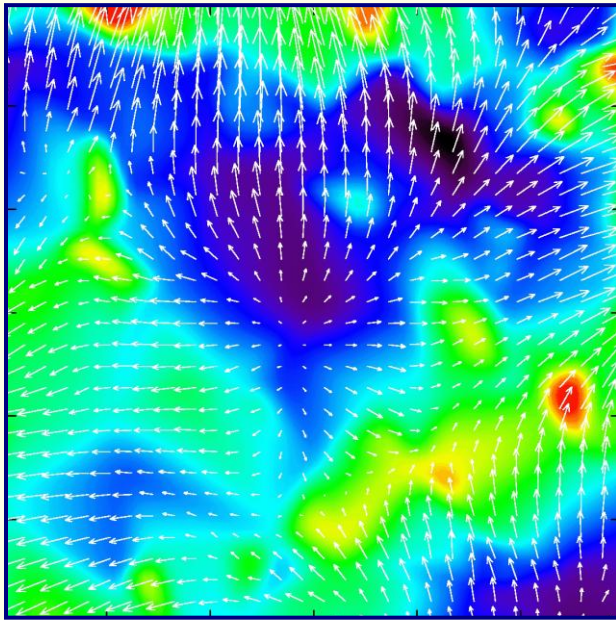
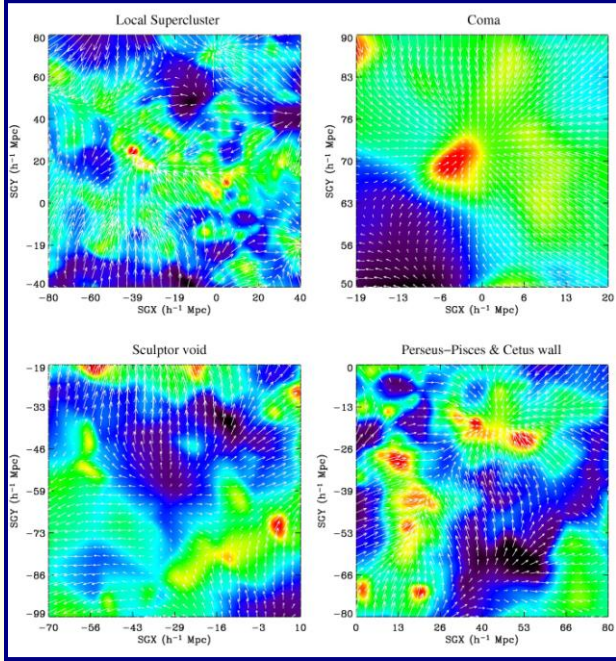
PSCz Flow Field

- Flow throughout the volume restored (assuming no shell-crossing)
- Restored velocities volume-weighted
- Shot-noise suppressed



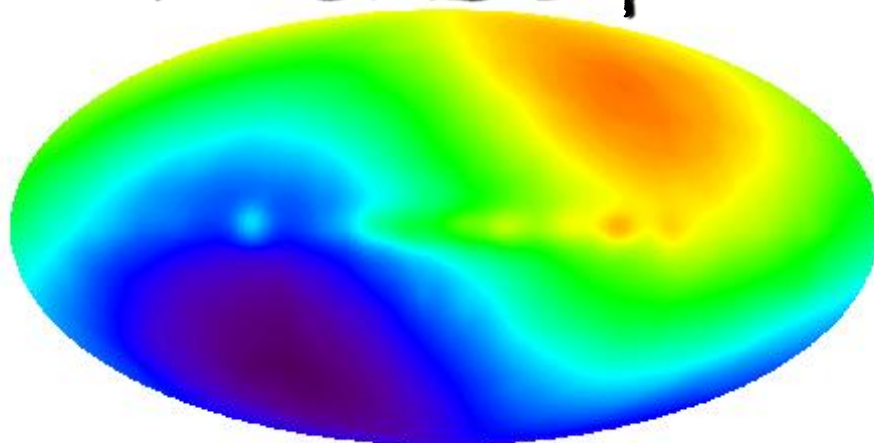
PSCz, velocity+density field





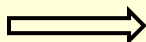
Cosmic Dipoles

The CMB Dipole



$$\Delta T = 3.36 \text{ mK}$$

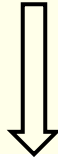
$$(l, b) = (264.3^\circ, 48.1^\circ)$$



$$v_{LG} = 627 \pm 22 \text{ km/s}$$

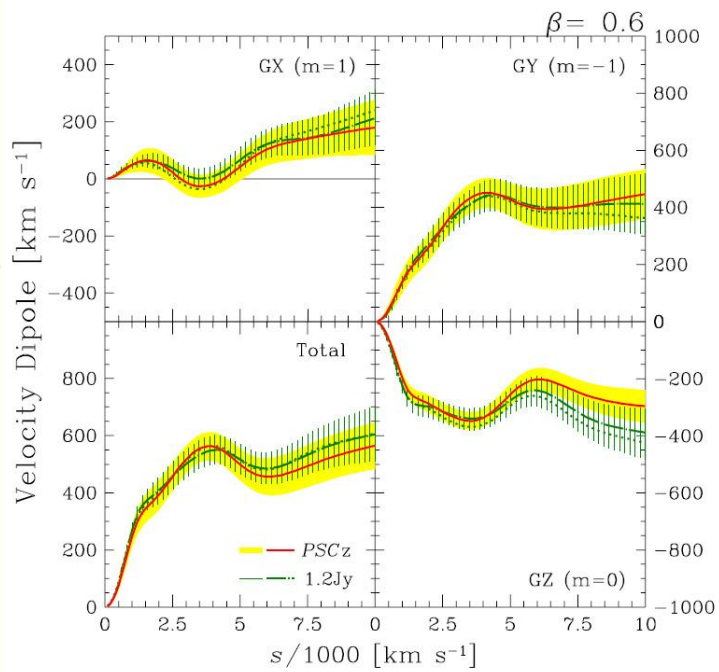
$$(l, b) = (276^\circ, 30^\circ)$$

$$\mathbf{v}_{\text{LG}} = \frac{H_0\beta}{4\pi} \int_r^\infty d^3\mathbf{r}' \delta_g(\mathbf{r}') \frac{\mathbf{r}' - \mathbf{r}}{|\mathbf{r}' - \mathbf{r}|^3}$$

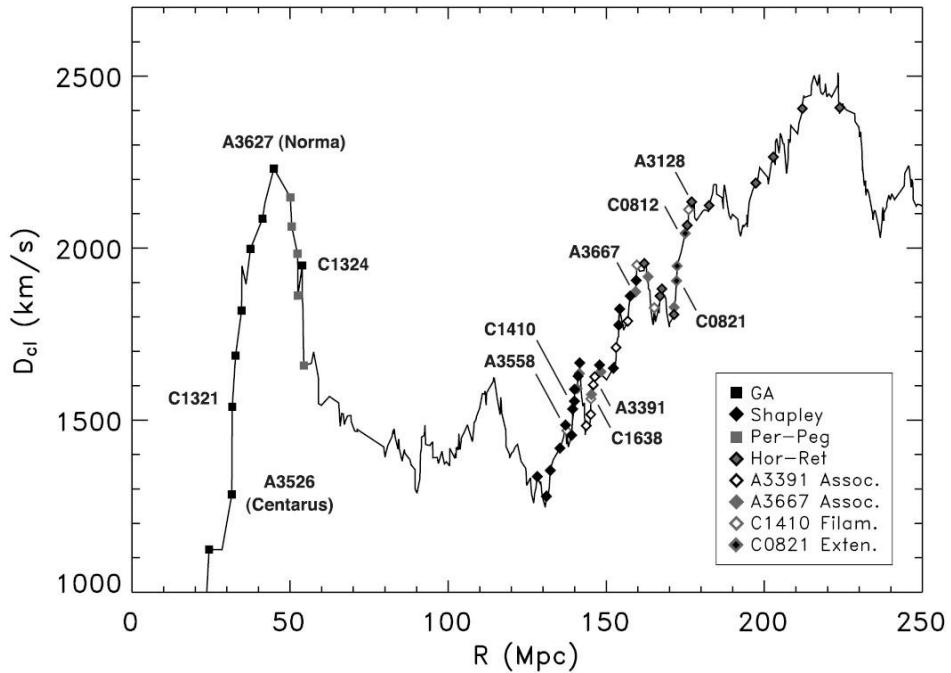
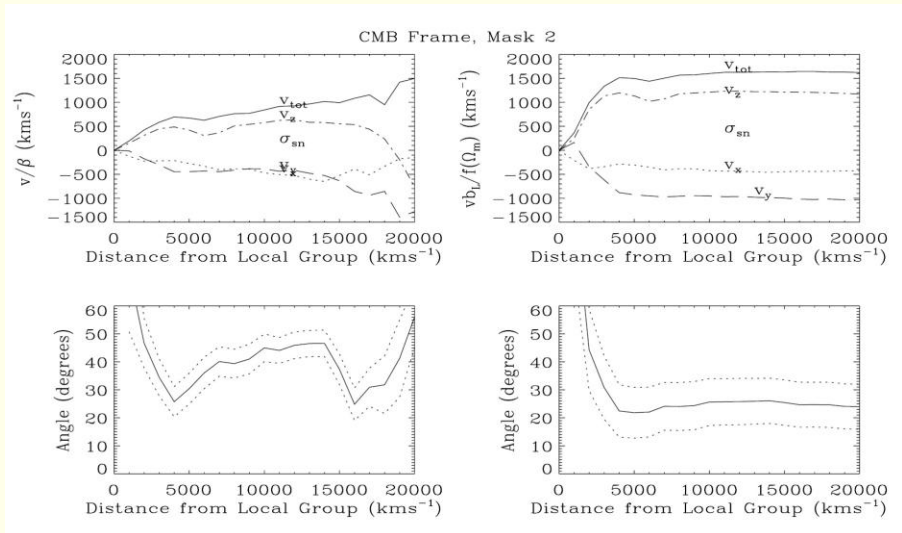


$$\mathbf{v}(\mathbf{r}) = \frac{H_0\beta}{4\pi\bar{n}} \sum_i^N \frac{w_i \hat{\mathbf{r}}_i}{r_i^2}$$

PSCz
dipole



2MASS dipole



2MASS survey

- 2MASS all-sky survey:
ground-based near-infrared survey whole sky,
J($1.2\ \mu\text{m}$), H($1.6\ \mu\text{m}$), K($2.2\ \mu\text{m}$)
- 2MASS extended source catalog (XSC):
1.5 million galaxies
- unbiased sample nearby galaxies
- photometric redshifts: depth in 2MASS maps,
“cosmic web” of (nearby) superclusters spanning
the entire sky.

courtesy:

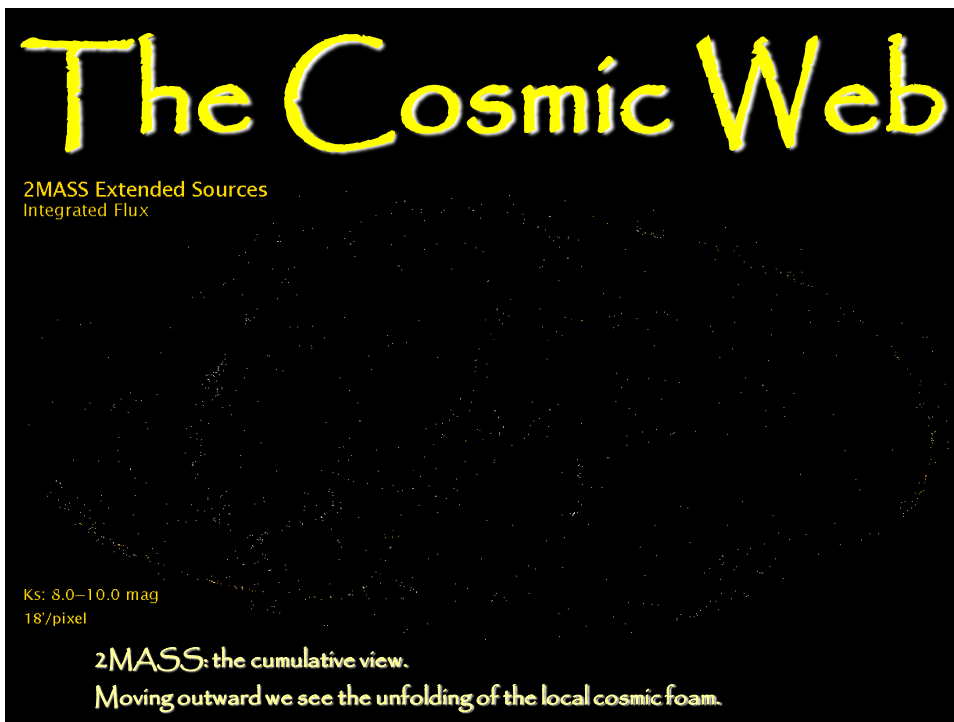
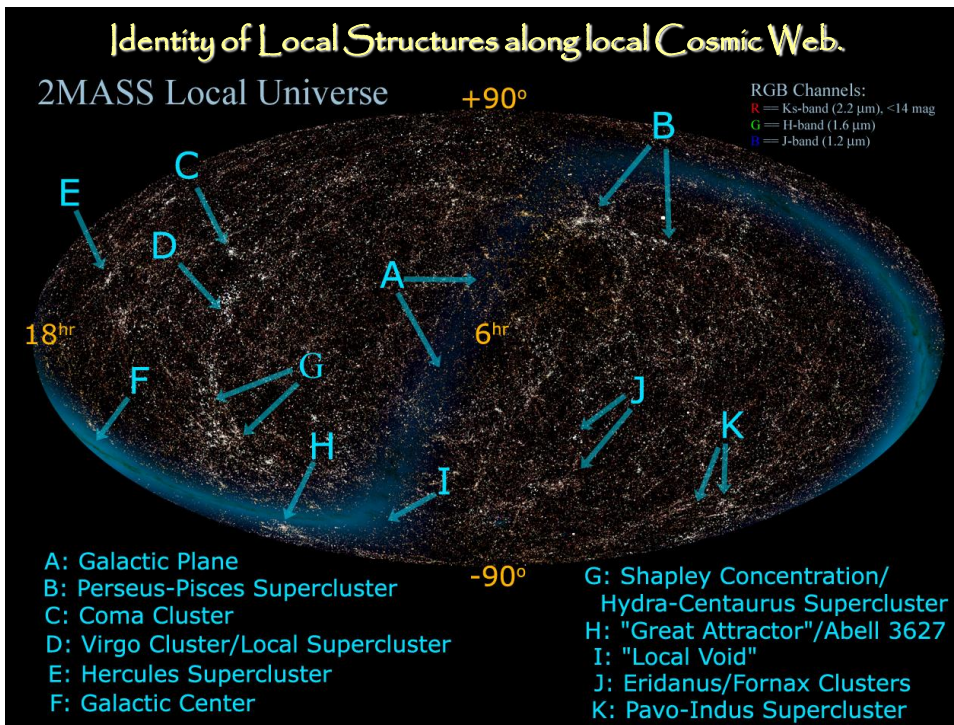
T. Jarrett

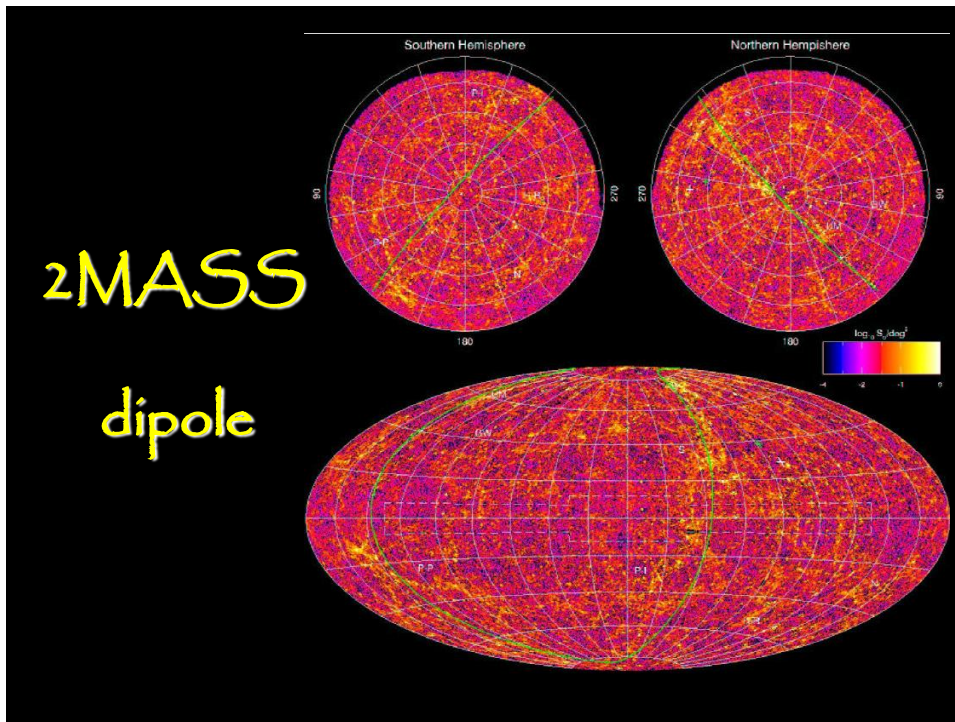
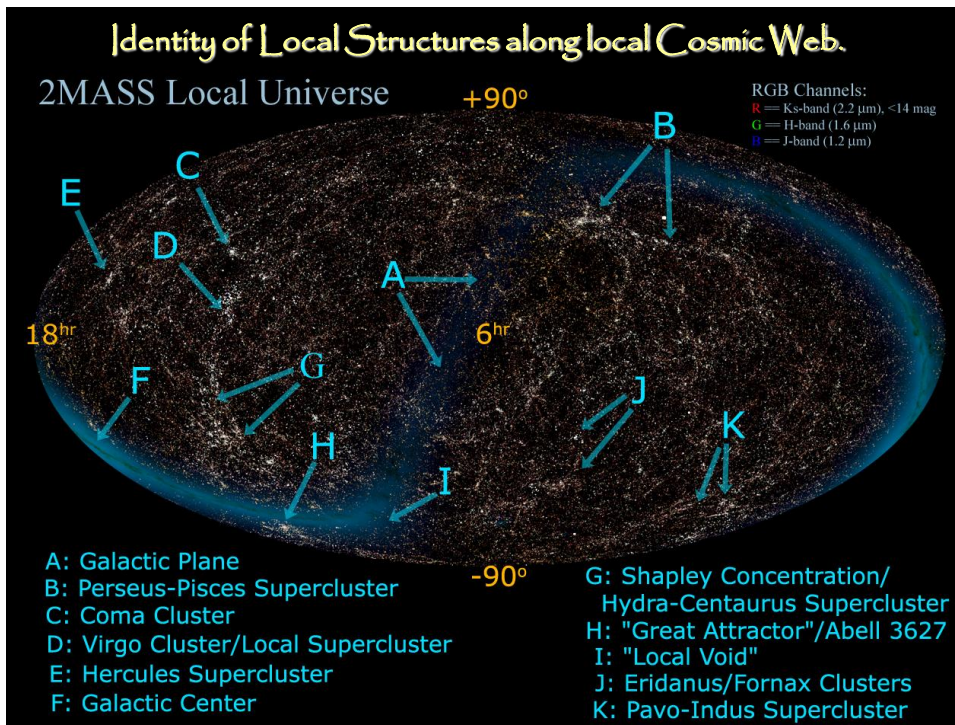
The Cosmic Web

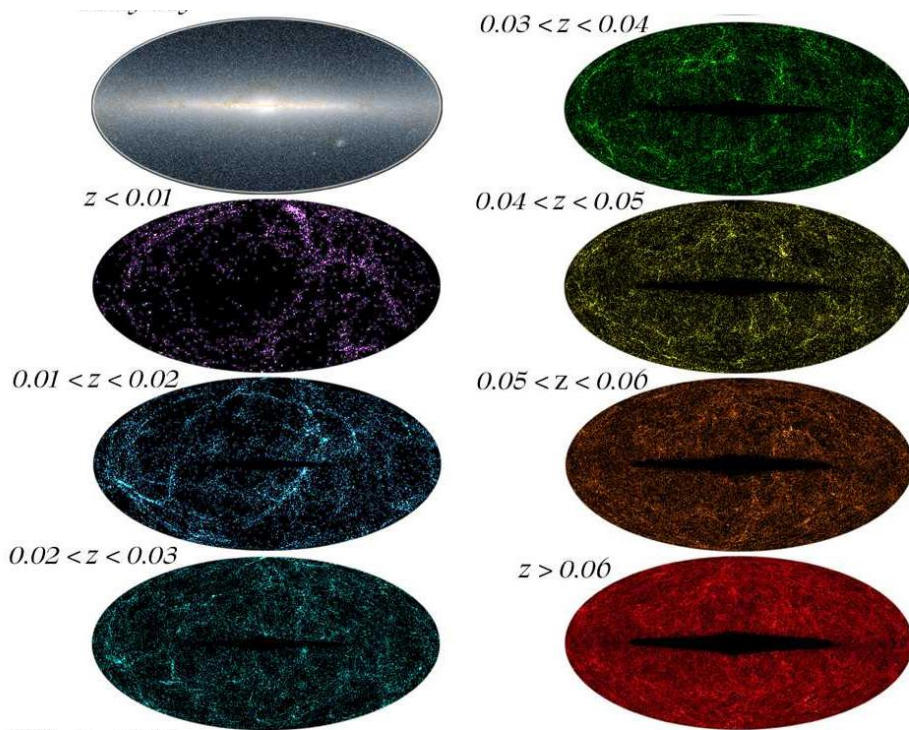
2MASS Extended Sources
Integrated Flux

Ks: 8.0–14.0 mag
18"/pixel

Looking around us we already see the unmistakable signatures of an intriguing
foamlike matter distribution in our immediate Cosmic Vicinity.







Redshift Space Distortions

